

Quantum Physical Relevance of the Einstein Tensor

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Taking quantum physics as well as large scale astronomical observations into account, a spacetime metric is introduced, such that the nonlinear part of the Einstein tensor contains effects of order \hbar .

1 Introduction

On large scales general relativity unveils that gravity appears as curvature of spacetime. Gravity and spacetime exist everywhere. It should therefore to be expected that doing quantum physics, learning more and more about the behavior of nature on small scales, would create an understanding of quantum physics in terms of curvature not only qualitatively but also quantitatively.

But this has not been the case so far in spite of the fact that the very notion of curvature itself, in the sense of obstruction to flatness, contains information on the internal structure of a curved manifold which in turn influences the local dynamics.

Nevertheless, with the Casimir and the Aharonov-Bohm effect two experimental observations are given in the quantum physical realm, which may give hints for an understanding of the role of curvature in quantum physics. As yet these experiments could not be explained in the context of general relativity. There are, however, new large scale observations from astronomy requiring analysis both along the lines of the Casimir and the Aharonov-Bohm effect as well as those of general relativity.

This text is organised as follows. After a short account of the Casimir and the Aharonov-Bohm effect, sec.2, we introduce two large scale observations: galaxy superclusters and the time development of the scale factor. Both are discussed with respect to their impact on quantum physics and general relativity, sec.3. In sec.4 we define values and formulas which we use in sec.5 to introduce a spacetime metric such that its curvature properties yield quantum physical orders of magnitudes. Consequences regarding the physical meaning of gravitational potentials in spacetime and the specific nonlinearity of general relativity are discussed.

The following convention is used. The metric tensor $g_{\mu\nu}$ is taken to have Lorentzian signature of type $(-,+,+,+)$, the Ricci tensor $R_{\mu\nu}$ is obtained from the Riemannian curvature tensor $R_{\mu\nu\rho\sigma}$ by contraction over the first and the fourth index, $R_{\mu\nu} = R_{\mu\nu\rho}^{\rho}$.

The Einstein field equations in geometric units are

$$G^{\mu\nu} = 8\pi T^{\mu\nu}. \quad (1)$$

Here $G^{\mu\nu}$ is the Einstein tensor,

$$G^{\mu\nu} = R^{\mu\nu} - \frac{1}{2}Rg^{\mu\nu}, \quad (2)$$

where R is the scalar curvature and $T^{\mu\nu}$ the energy momentum tensor.

2 On Casimir and Aharonov-Bohm effect

Calculating quantum fluctuations of the electromagnetic field between two perfectly conducting plates, placed parallel in vacuum, Casimir [1] predicted an attractive force F per unit area A between these two plates at a mutual distance d

$$\frac{F}{A} = -\frac{\pi^2 \hbar c}{240d^4}. \quad (3)$$

For a review of Casimir effect see monograph [2].

Remarkably, no constant characterizing an interaction is appearing in formula (3). This suggests the Casimir force to be of geometrical origin. In addition, the Casimir force per unit area is in fact a pressure. Due to general relativity and the Einstein field equations (1), any kind of pressure contributes to the curvature of spacetime. Thus an expression of Casimir force per unit area in terms of curvature should be expected.

Indeed, in geometric units c and the gravitational constant G are put equal to one and Planck's constant \hbar becomes about $\hbar^{gu} = 2.6 \times 10^{-66} \text{cm}^2$ [3]. Writing (3) for unit area $A = 1 \text{cm}^2$ and a distance $d = 1 \text{cm}$ we get

$$\frac{F}{A} = -0.04 \frac{\hbar^{gu}}{\text{cm}^4} = -0.04 \frac{2.6 \times 10^{-66}}{\text{cm}^2}. \quad (4)$$

This has the physical dimensions of intrinsic curvatures, like Gaussian, scalar or sectional curvatures for example.

Aharonov and Bohm [4] theoretically predicted that the vector potential of electrodynamics influences the interference of electrons in regions, where the magnetic field is effectively zero. For the history of experimental verification see monograph [5].

This shows that in quantum physics, in contrast to classical physics, dynamics is determined by potentials themselves and not by their derivatives alone; this way the gauge structure gets physically significant.

3 Large scale observations

Large scale astronomy delivers new observational results, which are relevant for general relativity as well as for quantum physics. Firstly, the improved precise measurement of large scale distances by supernovae Ia shows that the Hubble expansion is accelerated and thus a concrete value can be assigned to the time derivative of the Hubble function $H(t)$ at our epoche.

Secondly, the determination of the masses of the galaxy superclusters Great Attractor and Perseus Pisces leads to gravitational radii of about thousand parsecs - the greatest values to handle so far. These superclusters have a meaning for physics on Earth, which from a theoretical viewpoint is quite fundamental. Their gravitational potentials are constant in space through the entire solar system up to 10^{-11} , so they cannot yield observational effects in classical physics, because their spatial gradients are practically zero. The absolute values of these potentials, however, lie two or three orders of magnitude above that of the sun on Earth and well above the linear limit of general relativity discussed in [6].

The idea that the potential of the Great Attractor might lead to a gravitational Aharonov-Bohm effect has been developed in [7]. There the potential is added as an external one in the Schrödinger equation.

General relativity, however, is expressed mathematically in terms of differential geometry. Thus the question arises in which sense the nearly constant potentials of the Perseus Pisces as well as the Great Attractor cluster could be made accessible to differential geometry. The answer lies in focussing on time dependence. One is accustomed to observe the Hubble expansion through the recession of distant galaxies; but there is a possibility to observe the Hubble flow locally on the basis of the time dependence of exactly these gravitational potentials.

4 Formulas and values

In more detail, let $a(t)$ be the cosmological scale factor describing the average growth of physical distances. One expands $a(t)$ around the present time t_0 , into a Taylor series $a(t) = a(t_0) + \dot{a}(t_0)(t - t_0) + \frac{1}{2}\ddot{a}(t_0)(t - t_0)^2 + \dots$. Using the definition of the Hubble function

$$H(t) = \frac{\dot{a}(t)}{a(t)} \quad (5)$$

this gives up to second order

$$\frac{a(t)}{a(t_0)} = 1 + H_0(t - t_0) - \frac{q_0}{2}H_0^2(t - t_0)^2 \quad (6)$$

with $H_0 = H(t_0)$ and $q_0 = -\frac{\ddot{a}(t_0)}{a(t_0)}\frac{1}{H_0^2}$ today's Hubble function and deceleration parameter respectively. From observation one obtains $q_0 = -\frac{1}{2}$, which, using (5) and the definition of q_0 , implies

$$\dot{H}(t_0) = -\frac{1}{2}H_0^2. \quad (7)$$

In the calculations below the best present estimate of H_0 is chosen:

$$H_0 = \frac{1}{1.3 \times 10^{28} \text{cm}} \quad (8)$$

in geometric units.

The gravitational potential ϕ_{GA} of the Great Attractor supercluster in the solar system can be estimated to be [8]

$$\phi_{GA} = -\frac{r_g}{l} = -3.0 \times 10^{-5}, \quad (9)$$

analysing data of streaming motions of galaxy clusters in [9]. Hereby r_g is the gravitational radius in geometric units and l the distance from the center of the Great Attractor region to our Local Group.

Using mainly spectral methods, it is reported in [10] that the Perseus Pisces cluster happens to lie approximately at similar distances at opposite sides of the local group.

5 Quantum physical curvature

The Ansatz is the following. With z defining the redshift, the relation $z = H(t)l$ is used to gain the time dependence of the gravitational potentials ϕ_{GA} and ϕ_{PP} of the Great Attractor and the Perseus Pisces supercluster respectively. Now ϕ_{GA} becomes a function of t ,

$$\phi_{GA}(t) = -\frac{r_{gGA}}{z} H(t), \quad (10)$$

with $\phi_{GA}(t_0) = -3.0 \times 10^{-5}$. For simplicity ϕ_{PP} is assumed to have the same value as ϕ_{GA} and its time dependence is introduced in the same manner. Summing these two potentials and respecting Lorentzian signature a purely time dependent spacetime metric results

$$g_{\mu\nu} = \begin{pmatrix} -A(t) & 0 & 0 & 0 \\ 0 & B(t) & 0 & 0 \\ 0 & 0 & B(t) & 0 \\ 0 & 0 & 0 & B(t) \end{pmatrix} \quad (11)$$

with $A(t) = 1 + 2\phi(t)$ and $B(t) = 1 - 2\phi(t)$ and the indices dropped.

The component G^{00} of the Einstein tensor (2) of this metric can be calculated to be

$$G^{00} = -\frac{3}{4} \frac{\dot{B}(t)^2}{A(t)^2 B(t)^2}. \quad (12)$$

It is a well known fact that the G^{00} and G^{0i} components of the Einstein tensor contain only first derivatives of the metric with respect to time. Usually the Einstein equations of these components are taken as constraint equations.

However, evaluation of G^{00} with (10) and (7) at time t_0 and constant redshift z gives in leading order

$$G^{00}(t_0) = -\frac{3}{4}\phi(t_0)^2 H_0^2, \quad (13)$$

then using (8), (9) and dividing by 8π gives due to Einstein field equations (1)

$$T^{00} = \frac{G^{00}}{8\pi} = -0.06 \frac{2.6 \times 10^{-66}}{\text{cm}^2}, \quad (14)$$

which is nearly the value of the Casimir energy density per unit area and unit length given in (4).

Furthermore, G^{00} in (12) is due to the nonlinear part of the Einstein tensor, which is usually neglected in linear and Newtonian approximations of general relativity. But these approximations imply flat background and thus run against the notion of intrinsically curved spacetime.

6 Conclusions and outlook

Of course, this calculation is nonlocal in space, but 'locally' in full four-dimensional spacetime. One can insert the time dependent potentials in the post-Newtonian approximation developed in [12, 11]. Then nondiagonal terms would reflect the deviation of Perseus Pisces and Great Attractor being placed on a line. Nevertheless, G^{00} , but not G_{00} , would in leading order remain to be the same.

Also this Ansatz can be used to investigate differential topological properties of spacetime. This will be discussed in a future publication [13].

As a result we conclude, that Planck's constant appears to emerge from spacetime and represents the obstruction to flatness which is characteristic for the notion of curvature

Furthermore, as reported in [14] there are ambiguities regarding extent and nature of the Great Attractor which are lively discussed among astronomers. Also an influence of the Perseus Pisces cluster on the Local Group may be questionable, as stated for example in [15]. Perhaps this metric may be helpful to clarify the structure of the galaxy superclusters as well as the motion of the Local Group.

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References

- [1] H.B.G. Casimir, Proc. Sect. Sci. **51**, 793 (1948).
- [2] V.M. Mostepaneko and N.N. Trunov, *The Casimir effect and its applications* (Clarendon, Oxford, 1997).
- [3] R. Wald, *General Relativity* (Univ. of Chicago Press, Chicago, 1984).
- [4] Y. Aharonov and D. Bohm, Phys. Rev. **115**, 485 (1959).
- [5] A. Tonomura, *The Aharonov-Bohm effect* (Springer, Berlin, 1989).
- [6] C.W. Misner, K.S. Thorne, and J.A. Wheeler, *Gravitation* (W.H. Freeman and Company, New York, 1973).
- [7] D.V. Ahluwalia, Mod. Phys. Lett. A **13**, 1393 (1998), gr-qc/9805067.
- [8] I.R. Kenyon, Phys. Lett. B **237,2**, 274 (1990).
- [9] D. Lyndon-Bell *et al.*, Astrophys. J. **326**, 19 (1988).
- [10] R.C. Kraan-Korteweg, in: *From the Sun to the Great Attractor*, ed. by D. Page and J. Hirsch (Springer, Berlin, 2000) p. 301.
- [11] A. Einstein, L. Infeld, and B. Hoffman, Ann. Math. **39,1**, 65 (1938).
- [12] A. Eddington and G.L. Clark, Proc. Roy. Soc. A **166**, 465 (1938).
- [13] J. Lamey, Ph.D. thesis, to appear.
- [14] P.A. Woudt *et al.*, astro-ph/0410725v1.
- [15] J.L. Tonry *et al.*, Astrophys. J. **530**, 625 (2000).